Shortcomings of the R-Matrix Method for Treating Dielectronic Recombination

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By performing new radiation-damped R-matrix scattering calculations for the photorecombination of Fe¹⁷⁺ forming Fe¹⁶⁺, we demonstrate and discuss the difficulties and fundamental inaccuracies associated with the R-matrix method for treating dielectronic recombination (DR). Our new R-matrix results [1] significantly improve upon earlier R-matrix Results [2] for this ion. However, we show theoretically that all R-matrix methods are unable to account accurately for the phenomenon of radiative decay followed by autoionization. For Fe¹⁷⁺, we demonstrate numerically that this results in an overestimate of the DR cross section at the series limit, which tends to our analytically predicted amount of 40%. We further comment on the need for fine resonance resolution and the inclusion of radiation damping effects. Overall, slightly better agreement with experiment [3] is still found with the results of perturbative calculations, which are computationally more efficient than R-matrix calculations by more than two orders of magnitude.

References:

- [1] T. W. Gorczyca, N. R. Badnell, and D. W. Savin, Phys. Rev. A, in press (2002).
- [2] A. K. Pradhan, S. N. Nahar, and H. L. Zhang, Astrophys. J. 549, L268 (2001); H. L. Zhang,
 S. N. Nahar, and A. K. Pradhan, Phys. Rev. A 64, 032719 (2001).
- [3] D. W. Savin, T. Bartsch, M. H. Chen, S. M. Kahn, D. A. Liedahl, J. Linkemann, A. Müller, S. Schippers, M. Schmitt, D. Schwalm, and A. Wolf, Astrophys. J. 489, L115 (1997); D. W. Savin, S. M. Kahn, J. Linkeman, A. A. Saghiri, M. Schmitt, M. Grilser, R. Repnow, D. Schwalm, A. Wolf, T. Bartsch, C. Brandau, A. Hoffknecht, A. Müller, S. Schippers, M. H. Chen, and N. R. Badnell, Astrophys. J. Suppl. Ser. 123, 687 (1999).

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